

Interaction of ultra-energetic cosmic neutrinos with a thermal gas of relic neutrinos

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We use the formalism of finite-temperature field theory to study the interactions of ultra-high energy (UHE) cosmic neutrinos with the thermal background of relic neutrinos. From the imaginary part of the neutrino self-energy, calculated in terms of the Z boson propagator near the resonance, we derive general expressions for the UHE neutrino transmission probability. This allows us to take into account the thermal effects introduced by the momentum distribution of the relic neutrinos. We compare our results with the approximate expressions existing in the literature and discuss the influence of thermal effects on the absorption dips in the context of realistic UHE neutrino fluxes and favoured neutrino mass schemes.

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1. Introduction

The interaction of cosmic neutrinos at ultra-high energies (UHE ν) with the cosmological background of relic (anti)neutrinos (C ν B) has been proposed as a way of observing the C ν B, and a method to perform relic neutrino spectroscopy [1]. Provided adequate sensitivity and energy resolution of the detectors, the observation in the UHE ν flux of absorption lines associated with the resonant production of a Z boson ($\nu\bar{\nu} \rightarrow Z \rightarrow f\bar{f}$) could indeed allow an indirect determination of the absolute neutrino masses. The shape and depth of these absorption dips may also reflect features of the distribution of UHE ν sources and of their emission spectrum [2]. Most of the work in the literature describe the UHE ν -C ν B interactions assuming that relic neutrinos are at rest. However, effects of thermal motion in the C ν B (whose present temperature is $\approx 1.69 \times 10^{-4}$ eV) become relevant as soon as the momentum of the relic neutrinos gets comparable to their mass, and even before. To take this effect into account, we compute the dominant (resonant) contribution to the neutrino damping using the real-time formalism of finite-temperature field theory (FTFT), and investigate the modifications in the UHE ν transmission probability due to thermal effects [5].

2. Damping of UHE ν across the C ν B

For an UHE ν with four-momentum $k^\mu = (\mathcal{E}_k, \vec{K})$ and mass m_ν travelling across the C ν B, the equation of motion reads $(\not{k} - m_\nu - \Sigma)\psi = 0$, where the self-energy Σ embodies the effects of the medium. The corresponding dispersion relation is given by $\mathcal{E}_K = \mathcal{E}_r(K) - i\gamma(K)/2$. In our case, Σ is determined from a FTFT one-loop calculation carried out in terms of the (vacuum) Z boson propagator and the thermal propagator of the relic neutrinos. The last one depends on the functions $f_\nu(P)$ and $f_{\bar{\nu}}(P)$ which describe the momentum distributions of neutrinos (antineutrinos) in the thermal bath. These functions take the simple relativistic Fermi-Dirac form, $f_\nu(P) = f_{\bar{\nu}}(P) = 1/(e^{P/T_\nu} + 1)$, where T_ν is the temperature of the C ν B and we have neglected the chemical potential.

The damping factor γ governs the propagation of the UHE ν across the background of relic neutrinos and is directly related to the imaginary part of the self-energy, Σ_i [4]. In the approximation that the UHE ν are ultrarelativistic and we can neglect the background effects on their energy ($\mathcal{E}_r(K) \simeq K$), the damping can be written as (see [5] for the detailed calculation)

$$\gamma_{\nu\bar{\nu}}(K) = -\frac{1}{K} \text{Tr}(\not{k}\Sigma_i)|_{\mathcal{E}_r=K} = \int_0^\infty \frac{dP}{2\pi^2} P^2 f_{\bar{\nu}}(P) \sigma_{\nu\bar{\nu}}(P, K). \quad (2.1)$$

For $m_\nu \ll M_Z, K$ and neglecting terms of order Γ_Z^2/M_Z^2 , we have

$$\begin{aligned} \sigma_{\nu\bar{\nu}}(P, K) = & \frac{2\sqrt{2}G_F\Gamma_Z M_Z}{2KE_p} \left\{ 1 + \frac{M_Z^2}{4KP} \ln \left(\frac{4K^2(E_p + P)^2 - 4M_Z^2K(E_p + P) + M_Z^4}{4K^2(E_p - P)^2 - 4M_Z^2K(E_p - P) + M_Z^4} \right) \right. \\ & \left. + \frac{M_Z^3}{4K\Gamma_Z} \left[\arctan \left(\frac{2K(E_p + P) - M_Z^2}{\Gamma M_Z} \right) - \arctan \left(\frac{2K(E_p - P) - M_Z^2}{\Gamma M_Z} \right) \right] \right\}, \quad (2.2) \end{aligned}$$

where $E_p = \sqrt{P^2 + m_\nu^2}$ is the energy of the relic neutrino. Taking the limit of eq. (2.2) for $P \rightarrow 0$, one recovers the approximated cross-section used for relic neutrinos at rest, with the Z peak at the UHE ν "bare" resonance energy $K_{res} = M_Z^2/(2m_\nu)$. However, this approximation breaks down for

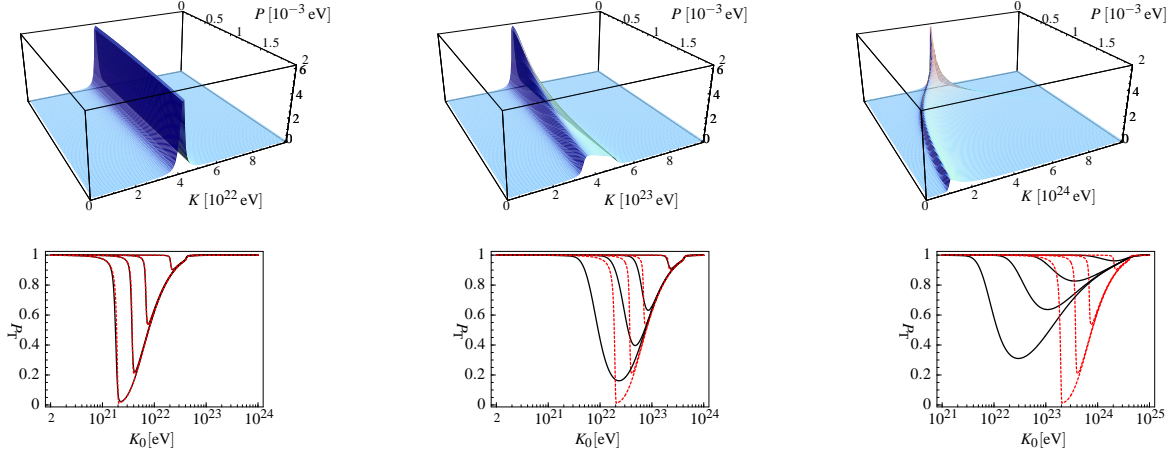


Figure 1: **Top :** cross-section $\sigma_{\nu\bar{\nu}}(P, K)$, in units of 10^{-31} cm^2 , as given by eq. (2.2), as a function of the energy of the incident neutrino, K , and of the relic neutrino momentum, P . From left to right, the panels correspond to a neutrino mass 10^{-1} , 10^{-2} , and 10^{-3} eV. **Bottom :** Transmission probability $P_T(K_0, z_s)$ as a function of the UHEv energy as detected on Earth, K_0 , for a source located at redshifts $z_s = 1, 5, 10, 20$ (from top to bottom in each panel) and for a neutrino mass $m_\nu = 10^{-1}, 10^{-2}, 10^{-3}$ eV (from left to right). The continued, black curves correspond to the full damping as given by eqs. (2.1) and (2.2), while the dotted (red) curves are for the approximation of relic neutrinos at rest.

small m_ν : Fig. 1 (top line) shows how the resonance peak in the $\nu\bar{\nu} \rightarrow Z$ cross-section broadens and shifts to lower UHEv energies as P increases. The transmission probability for an UHEv emitted at a redshift z_s to be detected on Earth with an energy K_0 is obtained by integrating the damping along the UHEv path, taking into account that both the UHEv energy and the CvB temperature are redshifted:

$$P_T(K_0, z_s) = \exp \left[- \int_0^{z_s} \frac{dz}{H(z)(1+z)} \gamma_{\nu\bar{\nu}}(K_0(1+z)) \right], \quad (2.3)$$

where $H = H_0 \sqrt{0.3(1+z)^3 + 0.7}$ is the Hubble factor. Fig. 1 (bottom line) shows that for $m_\nu/T_\nu \lesssim 10^2$ the absorption lines are also significantly broadened and shifted to lower energies, and that the effect increases with the distance travelled by the UHEv. This complicates the extraction of m_ν and z_s from the start- and endpoint of the absorption dip, which, in the approximation of relic neutrinos at rest, were respectively located at $K_0 = K_{res}/(1+z_s)$ and $K_0 = K_{res} = M_Z^2/(2m_\nu)$.

3. Absorption lines in the UHEv flux

To investigate this effect in a realistic context, we applied our calculation to a flux of UHEv

$$\mathcal{F}_\nu(K_0) = \frac{1}{4\pi} \int_0^\infty \frac{dz}{H(z)} P_T(K_0, z) \eta(z) J_\nu(K_0), \quad (3.1)$$

assuming a distribution of sources $\eta(z) = \eta_0 (1+z)^n \theta(z - z_{\min}) \theta(z_{\max} - z)$ with a common injection spectrum $J_\nu(K) = j_\nu K^{-\alpha} \theta(K_{\max} - K)$. The normalized flux then only depends on the difference of spectral indexes, $\alpha - n$, with typically $\alpha - n \approx 2$ for astrophysical (bottom-up) sources and $\alpha - n \approx 0$ for top-down processes [2]. For these two cases, we computed the normalized, all-flavour UHEv spectrum assuming some mass patterns compatible with the currently favoured three-neutrinos mass schemes [3]. Fig. 2 shows how thermal broadening affects the superposition of absorption lines and globally modifies the shape and extension of the dip.

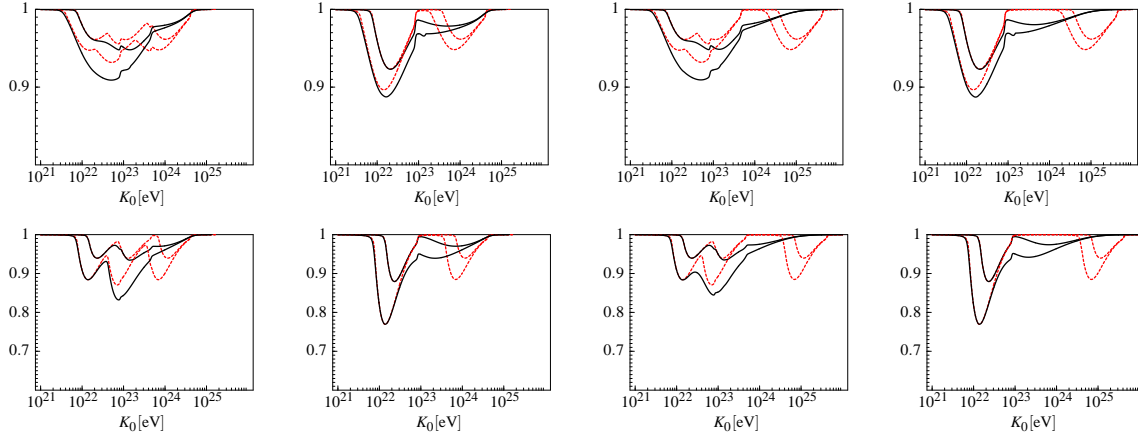


Figure 2: UHEV flux in presence of damping, \mathcal{F}_ν , normalized to the corresponding flux in absence of interactions. The top row is for $\alpha - n = 0$ and $z_s = 10, 20$, and the bottom row is for $\alpha - n = 2$ and $z = 5, 10$. From left to right, columns correspond to the following neutrino mass patterns: $\{10^{-3}, 9 \times 10^{-3}, 5 \times 10^{-2}\}$, $\{10^{-3}, 5 \times 10^{-2}, 5 \times 10^{-2}\}$, $\{10^{-4}, 9 \times 10^{-3}, 5 \times 10^{-2}\}$, $\{10^{-4}, 5 \times 10^{-2}, 5 \times 10^{-2}\}$ (all in eV). Colour code is as in fig. 1.

4. Conclusions

From the exploration of the parameter space currently allowed by astrophysical and cosmological constraints, we see that thermal effects do affect the transmission properties of UHEv across the CvB even in the regime of non-relativistic relic neutrinos. For most neutrino mass patterns, the extraction of the neutrino masses from the endpoint of the absorption lines is complicated by the broadening and merging of the dips, especially in normal hierarchical schemes (columns 1 and 3 in fig. 2) and for top-down-like injection spectra. Some information on the source distribution could still be provided by the onset energy and slope of the dip.

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References

- [1] T. J. Weiler, Phys. Rev. Lett. **49**, 234 (1982) and Astrophys. J. **285**, 495 (1984); E. Roulet, Phys. Rev. D **47**, 5247 (1993); H. Pas and T. J. Weiler, Phys. Rev. D **63** (2001) 113015; D. Fargion, P. G. De Sanctis Lucentini, M. Grossi, M. De Santis, and B. Mele, Mem. Soc. Ast. It. **73** (2002) 848.
- [2] B. Eberle, A. Ringwald, L. Song, and T. J. Weiler, Phys. Rev. D **70**, 023007 (2004); G. Barenboim, O. Mena Requejo, and C. Quigg, Phys. Rev. D **71** (2005) 083002.
- [3] J. F. Beacom and N. F. Bell, Phys. Rev. D **65** (2002) 113009.
- [4] J. C. D’Olivo and J. F. Nieves, Phys. Rev. D **52** (1995) 2987.
- [5] J. C. D’Olivo, L. Nellen, S. Sahu and V. Van Elewyck, arXiv:astro-ph/0507333, to appear in Astropart. Phys.